

GRAVITATIONAL WAVES, EXTREME ASTROPHYSICS, AND FUNDAMENTAL PHYSICS WITH PRECISION PULSAR TIMING

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FOR THE NANOGrAV COLLABORATION
(*The North American Nanohertz Observatory for Gravitational Waves*)

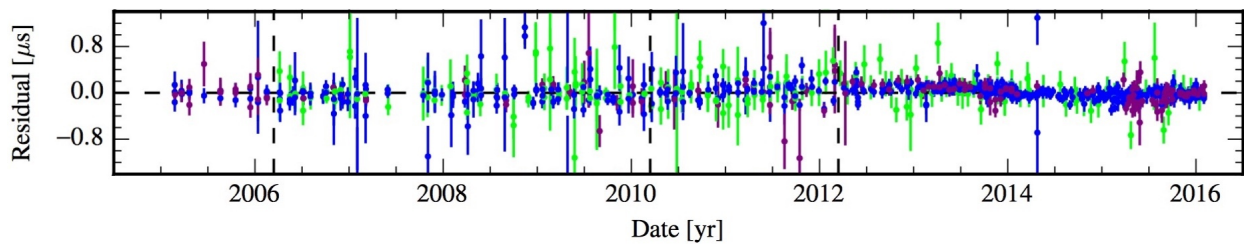


Figure 1: Over a decade of timing residuals from millisecond pulsar J1713+0747 with the Arecibo Observatory and Green Bank Telescope at 820 MHz (green), 1.4 GHz (blue), and 2.1 GHz (magenta) [1]. The RMS residual is 116 ns, demonstrating the remarkable stability of millisecond pulsars over long timespans.

This is one of five core white papers written by members of the NANOGrav Collaboration.

Supermassive Black-hole Demographics & Environments with Pulsar Timing Arrays, S. Taylor et al.

Fundamental Physics with Radio Millisecond Pulsars, E. Fonseca et al.

Physics Beyond the Standard Model with Pulsar Timing Arrays, X. Siemens et al.

Multi-messenger Astrophysics with Pulsar Timing Arrays, L.Z. Kelley, M. Charisi et al.

OTHER RELATED SCIENCE WHITEPAPERS

Gravitational-Wave Astronomy in the 2020s and Beyond: A view across the gravitational wave spectrum, the Gravitational Waves International Committee

The Virtues of Time and Cadence for Pulsars and Fast Transients, R. Lynch et al.

Twelve Decades: Probing the ISM from kiloparsec to sub-AU scales, D. Stinebring et al.

- Thematic Areas:**
- Planetary Systems
 - Star and Planet Formation
 - Formation and Evolution of Compact Objects
 - Cosmology and Fundamental Physics
 - Stars and Stellar Evolution
 - Resolved Stellar Populations and their Environments
 - Galaxy Evolution
 - Multi-Messenger Astronomy and Astrophysics

1 Probing the Universe with Precision Timing

Precision pulsar timing at the level of tens to hundreds of nanoseconds allows detection of nanohertz gravitational waves (GWs) from supermassive binary black holes (SMBBHs) at the cores of merging galaxies and, potentially, from exotic sources such as cosmic strings. The same timing data used for GW astronomy also yield precision masses of neutron stars orbiting other compact objects, constraints on the equation of state of nuclear matter, and precision tests of General Relativity, the Strong Equivalence Principle, and alternative theories of gravity. Timing can also lead to stringent constraints on the photon mass and on changes in fundamental constants and could reveal low mass objects (rogue planets, dark matter clumps) that traverse pulsar lines of sight. Data sets also allow modeling of the density, magnetic field, and turbulence in the interstellar plasma. Roughly 100 millisecond pulsars (MSPs) are currently being timed at \sim GHz frequencies using the largest radio telescopes in the world. These data sets currently represent \sim 1000 pulsar-years and will increase dramatically over the next decade. These topics are presented as a program of key science with flowdown technical requirements for achieving the science.

2 Pulsar Timing

The time of arrival of astrophysical pulses is determined by the path of the pulse, and the intervening magnetized plasma that invariable exists between source and observer.

The precision of a time of arrival (TOA) depends on the pulse width and on the signal-to-noise ratio, thus favoring objects that emit bright, narrow pulses and observations using large telescopes with sensitive, wide-bandwidth receivers. At present the best TOA precisions of tens of ns are achieved with 100-m class or larger telescopes at \sim GHz frequencies and GHz bandwidths (see Fig. 1 and [1]). These precisions can be improved for a large number of MSPs using a variety of methods we discuss below.

GWs and other phenomena produce TOA departures from a comprehensive model that is constructed for each pulsar. These models account for all known effects, including the spin, spindown, position, proper motion, and parallax of the pulsar, along with those extrinsic to the pulsar, such as any binary motion, the movement of our local reference frame, and frequency-dependent delays from the interstellar plasma. GW detection requires sophisticated noise models that include astrophysical effects and measurement errors. [2]. The primary astrophysical effects are pulse jitter intrinsic to the pulsar [3, 4], low-level noise in the pulsar spin [5, 6], and imperfect removal of chromatic interstellar propagation effects [7, 8, 9, 10]. The accuracy of timing models is quantified by the RMS TOA residuals (measured minus predicted arrival times), which include TOA errors and other contributions. RMS values $\sigma_{\mathcal{R}} < 100$ ns over timespans of decades have been achieved [1, 11].

Since the first MSP discovery in 1982, radio surveys have revealed about 300 MSPs in our Galaxy and another 150 in globular clusters [12]. Ensembles of Galactic MSPs, or pulsar timing arrays (PTAs), are being observed around the world with multiple telescopes to search for the TOA perturbations expected from nanohertz GWs. Roughly 100 MSPs are currently included in this effort although this number continues to grow as new, bright MSPs are found. These datasets also allow a wide variety of other fundamental measurements, including the masses of compact objects (neutron stars, white dwarfs) in binary systems, precision tests of General Relativity, limits on the photon mass, and constraints on dark matter clumps and cosmic strings.

In the next decade, increased access to large telescopes along with algorithmic advances will lead to discoveries in all of these science areas.

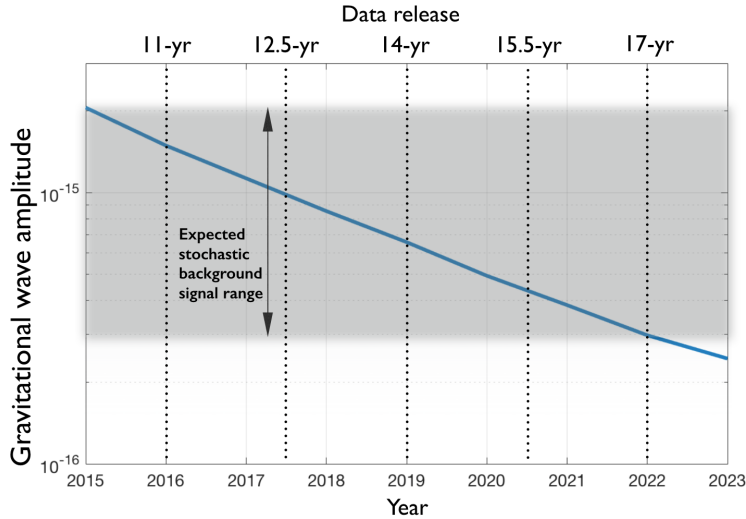


Figure 2: Detectable amplitudes of a stochastic background of GWs produced by SMBBH mergers as a function of time [1, 19]. The blue curve corresponds to the 95% upper limit expected to be set by NANOGrav in the absence of a GW signal. The shaded region corresponds to the $1\text{-}\sigma$ range of recent background estimates produced using a range of galaxy merger rates and empirical black hole-host relations [20, 21, 22]. The dotted lines show the cut-off times of previous and planned NANOGrav data releases.

3 Detection and Characterization of Nanohertz Gravitational Waves

To exploit MSPs as astrophysical clocks, they have been assembled into Galactic-scale detectors (PTAs) of long-wavelength gravitational waves. PTAs can detect strains $h \sim \sigma_{\mathcal{R}}/T$ at frequencies where they are most sensitive, $f_{\text{GW}} \sim 1/T$, where T is the total timespan of observations. Decade-long timing data sets with $\sigma_{\mathcal{R}} \sim 100$ ns therefore probe strains of 10^{-15} at frequencies $\sim 10^{-8}$ to 10^{-9} Hz. Detection and identification of GWs is aided by the angular correlation of timing residuals between pairs of pulsars. The angular correlation function is known for General Relativity [13, 14, 15, 16] and for alternate theories of gravity [13, 17, 18].

Galaxy mergers form the backbone of current structure-formation models [23]. As a byproduct, their central supermassive black holes form SMBBHs that eventually merge due to GW emission in the nanohertz regime. These comprise the dominant expected sources of low-frequency GWs. They may be detected individually as continuous wave sources or as burst sources from the memory effect at the time of merger of the SMBBH. However, the collective *stochastic background* from the incoherent superposition of GWs from the cosmic history of black hole mergers is likely to be detected first with dimensionless strain $h \sim 10^{-15}$. The stochastic background is expected to have a steep spectrum that increases at lower frequencies. TOA perturbations therefore grow $\propto T^{5/3}$ as longer data spans sample increasingly lower frequencies [24]. Long, continuous data sets obtained with stable instrumentation are a fundamental requirement for nHz GW science.

With modest projected increases in current sensitivity in the near term as additional MSPs are included, the stochastic background is expected to be detected within the next several years (Fig. 2) and GWs from individual SMBH binaries before the end of the next decade [25] (Fig. 3).

Other possible GW sources include cosmic strings and primordial GWs (see white paper *Physics Beyond the Standard Model with Pulsar Timing Arrays*, Siemens et al.). The most sensitive limits on the stochastic GW background already inform models for galaxy evolution, disfavoring those with large black hole to galactic bulge mass ratios and placing

constraints on the astrophysical processes involved in galaxy mergers, demonstrating that they are not solely driven by gravitational wave emission but also through processes such as stellar hardening, eccentricity, and gas dynamics [19, 26] (see white paper *Supermassive Black-hole Demographics & Environments with Pulsar Timing Arrays*, Taylor et al.). They also set the most stringent limits on the properties of cosmic strings [19] and can be used to place sensitive limits on individual SMBBH sources in local group galaxies [27].

PTAs are also sensitive to Galactic sources of GWs and to small masses that can perturb the location of the pulsar or the solar system or that perturb arrival times directly through gravitational lensing or the Shapiro delay [28, 29, 30, 31]. Bursts with memory are potentially detectable along lines of sight to pulsars close to any merging binaries, such as in the Galactic center or in globular clusters [32, 33, 34].

4 Understanding the Physics of Pulsar Spins

The rapid spins of MSPs comprise the astrophysical clocks that underlie their use for gravitational wave detection. Most pulsars, including some MSPs, show spin fluctuations that contribute red noise that limits GW sensitivity at very low ($\lesssim 0.1$ cy yr $^{-1}$) frequencies. Red noise is much weaker in MSPs compared to other pulsars, but its physics is not well understood. Longer data sets on more MSPs will allow better empirical characterization that will further optimize GW detection. Development of a better scaling law of the RMS red noise vs. spin rate, magnetic field, and other parameters will allow better forecasting of timing quality for selection of MSPs for inclusion in PTAs. Masses of MSPs, which vary from ~ 1.3 to $2 M_{\odot}$, may be one of the ‘hidden’ variables in the scaling law [5]. Increasing the MSP mass measurement sample to 100 or more is key to exploring this further.

5 Fundamental Physics

Timing of pulsars in double neutron star (DNS) systems has famously provided tests of General Relativity to better than 0.01% precision [35, 36]. More extreme binary systems remain to be discovered (e.g. DNS systems with sub-hour orbits, stellar mass black hole-pulsar systems, pulsars orbiting Sgr A*); they will provide fundamentally stronger constraints on alternative theories of gravity and on the environment of the Galactic center. Post-Newtonian terms in relativistic binaries also yield precision masses of neutron stars and their companions that have ruled out many equations of state for supranuclear matter [37, 38]. Timing observations of a triple system (a pulsar with two white dwarf companions) have placed the best constraints yet on the Strong Equivalence Principle [39]. Alternative theories of gravity also impact the properties of gravitational waves and their timing perturbations. The angular correlation, for example, is altered if there are more than two GW polarizations [17] (see whitepaper *Fundamental Physics with Radio Millisecond Pulsars*, Fonseca et al.).

6 Exotic Physics

Numerous low-mass objects in the Galaxy can produce detectable gravitational perturbations on an individual pulsar or on the Solar System and thus a collective effect on all lines of sight. Gravitational lensing from passage close to pulsar line of sight is also possible, with the small optical depth requiring many pulsar-years of timing measurements. Dark matter clumps on mass scales of 10^{-11} to $10^{-8} M_{\odot}$ [40] are potentially detectable after a few thousand pulsar-years of high-precision ($\sigma_{\mathcal{R}} \sim 10s$ of ns) timing [41], achievable in a long-term project.

Warm or “fuzzy” dark matter [42, 43] may also cause oscillations of the Galactic gravitational potential, which would manifest as low-frequency gravitational waves; PTA data currently

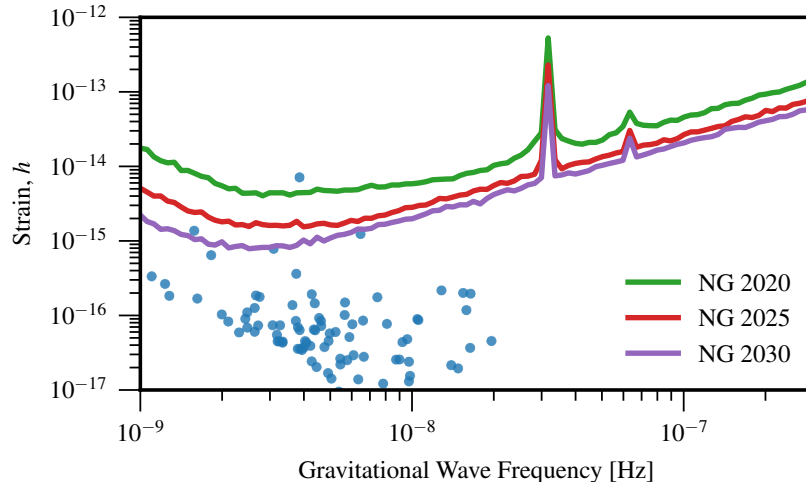


Figure 3: Example realization of the local GW sky, with nanohertz sources as blue dots. The curves are 95% detection probability for NANOGrav in 2020, 2025, and 2030, with a 3×10^{-3} false alarm rate [25]. Averaging over realizations, one to several individual SMBBHs should be detectable within the next decade.

sets the tightest constraints on the presence of ultralight bosons [44]. Timing of binary MSPs can also be used to place constraints on a variable gravitational constant and MSP spin-down measurements may also be used to constrain the mass of the photon (see white paper *Physics Beyond the Standard Model with Pulsar Timing Arrays*, Siemens et al.).

7 Science Requirements

The full science program requires suitable pulsars, including those currently known and the anticipated larger number of new MSPs and relativistic binaries that will be discovered next decade. The program also requires sustained observations on stable instrumentation. In particular, PTA science requires observations of 100 to 200 MSPs at regular cadence over time spans of decades. Uniform sky coverage is required to probe the detailed spectrum and isotropy of the stochastic background and to detect individual sources. Uniform GW sensitivity over the sky necessarily requires multiple telescopes for the timing program and it is essential to have overlapping coverage on a few objects for calibration purposes. *This will require increased access to high-sensitivity telescopes beginning as soon as possible.*

Improving Sensitivity to GWs and Post-Newtonian Effects:

Longer Data Sets. Precision MSP timing data sets using modern instrumentation currently span one to two decades, but increasingly longer data sets will have transformative effects on PTA science. The signals from the stochastic GW background and from bursts with memory grow as $\sim T^{5/3}$ and T , respectively. A crucial point is that errors in the solar system ephemeris decouple from GW signals for data sets longer than the orbital periods of Jupiter and Saturn (which are comparable to current data set lengths), providing greater sensitivity to GWs. Further benefits of long datasets include increased sensitivity to small bodies in the Solar System [45] or Milky Way, and to dark matter substructure [46].

Greater collecting area. Timing precision of the majority of MSPs in PTAs is limited by radiometer noise [1]. Noise modeling indicates that telescopes with sensitivities up to twice that of Arecibo would significantly improve TOAs for most currently timed pulsars and would allow many more objects to be used.

Higher observational cadences and wider bandwidths. Higher cadence (i.e. daily–weekly) and wide-bandwidth observations will increase TOA precision via better mitigation of chromatic interstellar delays [1]. Faster sampling also increases the probability of identifying rapid gravitational effects from dark matter clumps.

Larger numbers of MSPs. A PTA’s sensitivity to GWs increases linearly with the number of MSPs. For this reason, and to achieve uniform GW sensitivity across the sky, it is critical to increase the number of well-timed MSPs to at least 200 for detection and full characterization of the stochastic GW background (spectrum and isotropy). Gaps in sky coverage yield nonuniform GW sensitivity by a factor of about seven [27]. This can be rectified by targeting MSP surveys in particular sky locations. Continued surveys are also critical for discovering extreme relativistic binary systems.

Improved calibration and algorithms. TOA precision is limited at the tens of ns level by instrumentation, such as the quality of analog-to-digital converters and post-calibration polarization purity, which alter measured pulse shapes and the resulting TOAs. New instrumentation and polarization calibration methods can provide significant improvement. Timing precision will also benefit greatly from improved algorithms for radio frequency interference rejection and advances in the correction of ISM delays.

Solar System Ephemerides. Timing data are analyzed in the frame of the Solar System’s barycenter. Barycentric referencing is achieved by removing appropriate light-travel times estimated from JPL-provided ephemerides. Their present accuracy is currently a limiting factor in the detection of the stochastic GW background because systematic effects appear at the orbital frequency of Jupiter when different ephemerides are compared [19]. Continued space missions will improve the ephemeris and longer timing data sets will provide independent constraints on barycentric corrections.

Better distance measurements. Interferometric (VLBI) distances are necessary for distinguishing period derivatives due to Galactic acceleration from those induced by relativistic effects. This is critical for tests of GR using MSP binaries. Furthermore, if pulsar distances can be determined to within a GW wavelength (i.e. \sim parsecs) the “pulsar term” can be included in CW searches, enabling Galactic-scale GW interferometry [47] and increasing sensitivity by roughly a factor of $\sqrt{2}$ and dramatically improving localization capability. Distance estimates to binary MSPs can also be improved by observing the optical emission of their companions with e.g. Gaia [48].

8 Prospects for 2020-2030

New discoveries in the next decade will provide more extreme relativistic binaries and additional, high-quality MSPs for gravitational wave detection. Large scale surveys for pulsars require thousands of hours on large radio telescopes and deployment of additional machine-learning algorithms. Sustained timing programs on these objects will yield more stringent tests of theories of gravity and detection and characterization of low-frequency gravitational waves. These GWs will provide unique information on the mergers of supermassive binary black holes and on cosmic strings and other exotic potential sources of GWs.

Fulfillment of NANOGrav’s GW program [1] will be coordinated through the International Pulsar Timing Array (IPTA) consortium [49, 50], which pools data from around the world to provide greater GW sensitivity. These efforts will allow us to fully exploit the promise of low-frequency GW astronomy and multi-messenger astrophysics to probe galaxy formation and evolution through cosmic time.

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