Astro2020 Science White Paper

Inflation and Dark Energy from spectroscopy at $z > 2$

Thematic Areas: ☐ Planetary Systems ☐ Star and Planet Formation
☐ Formation and Evolution of Compact Objects ☑ Cosmology and Fundamental Physics
☐ Stars and Stellar Evolution ☐ Resolved Stellar Populations and their Environments
☐ Galaxy Evolution ☐ Multi-Messenger Astronomy and Astrophysics

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The expansion of the Universe is understood to have accelerated during two epochs: in its very first moments during a period of 'Inflation' and much more recently, at $z < 1$, when an unknown element known as Dark Energy drives cosmic acceleration. The undiscovered mechanism behind these two epochs represent some of the most important open problems in fundamental physics.

Most of the processes involved during Inflation impact observations on the very largest spatial scales [1, 2]. Traditionally, these have been accessed through observations of the Cosmic Microwave Background (CMB). While very powerful, the CMB originates from a 2D surface and the finite number of modes that it contains will largely be measured by experiments over the next decade[1]. Observations of large 3D volumes with large-scale structure (LSS) access similar scales and will dramatically increase the number of available modes. For example, LSS observations in the range $2 \lesssim z \lesssim 5$ can more than triple the volume surveyed at $z \lesssim 2$, and, together with the sufficiently high galaxy number in this interval, strongly motivates a future spectroscopic survey that exploits this opportunity. In addition, tomography allows a mapping of the growth of structure with redshift, which provides robust constraints on Dark Energy and neutrino masses, while relaxing restrictive assumptions such as a power-law primordial power spectrum [7].

Finally, cross-correlation with external tracers, such as CMB lensing, Intensity Mapping or the Lyman-α forest, immunises the constraints to the systematics that make measurement challenging and further improves the precision through 'sample variance cancellation' [8, 9, 10] and degeneracy breaking.

1 Science Case

Inflation Simple theories of inflation, involving a single non-interacting field, predict that the primordial fluctuations are extremely close to Gaussian distributed [11, 12]. However, very large classes of inflationary models produce levels of non-Gaussianity that are detectable by the next generation of spectroscopic surveys [1]. Measurements of primordial non-Gaussianity probe the dynamics and field content of the very early Universe, at energy scales far above particle colliders. Deviations from Gaussianity leave a particular imprint on the galaxy three-point correlation function or bispectrum [13] (and of the CMB), and can also produce a characteristic scale-dependence in the galaxy bias [14]. Depending on the physical process responsible for these deviations from Gaussianity, different configurations in the three-point function are generated. These are typically described by a number of dimensionless parameters, f_{NL} [15], and common examples include the local, equilateral and orthogonal types. The local type is generically produced in multi-field inflation, while the equilateral type often indicates self-interaction of the inflaton.

Pushing the observational frontier to the threshold typically expected from 'non-minimal' inflation ($f_{NL} \gtrsim 1$, see [2]) provides a compelling opportunity for future large-scale structure surveys. In summary, capturing the full picture of inflation requires measuring primordial non-Gaussianity to an unprecedented level, complementing the search for primordial gravitational waves and informing us about the Universe's first moments.

[1]Cosmologically relevant modes of CMB temperature have been measured to the cosmic-variance limit by Planck [3] and upcoming or proposed experiments will achieve the same for polarization [4, 5, 6].

Dark Energy A large number of theories have been put forward to explain the late time cosmic acceleration. They range from a cosmological constant, to some dynamical forms of Dark Energy or modification to General Relativity on large scales [16, 17]. By mapping expansion and growth at $z > 1.5$ – deep into matter domination – we can ease parameter degeneracies, better constrain potential theories of Dark Energy, and test posited modifications to General Relativity, e.g. by comparing measurements of growth to the amplitude of gravitational lensing of the CMB.

Curvature A measurement of the global value of the Universe's curvature can potentially have important implications for Inflation. Slow-roll eternal inflation predicts $|\Omega_K| < 10^{-4}$, while false-vacuum models would be ruled out by a measurement of $\Omega_K < -10^{-4}$ [18, 19]. Moreover, the current bound $\Omega_K < 2 \times 10^{-3}$ [3] relies on the strong assumption that Dark Energy is a cosmological constant. If this is relaxed, large degeneracies with the time evolution of Dark Energy arise, significantly degrading the constraints on both. Measurements at high redshift can break this degeneracy and, at the same time, approach the threshold $\sigma(\Omega_K) \approx 10^{-4}$ that is crucial for a better understanding of Inflation [20].

Neutrino Masses Massive neutrinos suppress the growth of structure on small scales in a time-dependent manner [21]. Measuring the amplitude of structure over a long lever-arm in redshift, $z \sim 0 - 5$, better constrains the neutrino masses and breaks important degeneracies with the time evolution of Dark Energy and the primordial power spectrum [22, 23].

1.1 High-z Lyman-break galaxies and Lyman-α emitters

Lyman-break galaxies are young, star forming galaxies that comprise the majority population at $z > 1.5$. Their characteristic spectral energy density exhibits a sharp drop in the optical flux blue-wards of the redshifted Lyman limit, $(1 + z) \times 912$Å, due to absorption by neutral hydrogen, in an otherwise shallow F_ν spectrum. As such, they are efficiently selected with a search for galaxies bright in a detection band, m_{UV} – chosen to correspond to the rest-frame UV for ease – but otherwise undetected in all bluer filters; See Refs. [24, 25] for reviews. In this manner, convenient target populations (BX, u-dropouts, g-dropouts and r-dropouts) spanning $\Delta z \simeq 1.0$ at $z \simeq 2, 3, 4$ and 5 are obtained by enforcing these criteria for increasingly red detection bands; Selection on photometric redshift largely yields the same ends [26, 27].

While of great interest for providing very large populations at high redshift, to achieve the necessary spectroscopic success rate in a baseline exposure typically requires refinement to those with significant Lyman-α emission (LAEs). This is traditionally achieved with narrow-band selection, but large volumes and sufficient depth are not obtainable in this manner. Accepting some degree of increased contamination or lower completeness, broad-band selection based on the bluer continua of strong emitters has been shown to provide very encouraging results [28, 29, 30]. Alternatively, one may limit oneself to only the brightest galaxies, for which secure absorption line redshifts are also possible.

1.2 Survey strategy

We identify two galaxy surveys that we use as a baseline for forecasts of an airmass-limited 14,000 square degree survey. Following Ref. [10], we first consider the idealised $m_{UV} = 24.5$

sample in Table 1. This informs what conclusions may ultimately be drawn for this science case with minimal assumptions on the required facilities and survey details.

Conversely, assuming a next generation survey speed, we posit a fiducial survey to approximate the properties shown in Table 2 – assuming completion of LSST Year 10 by first light; Interim LSST data may suffice depending on the cadence strategy adopted.

z	$n(z)$ $[10^{-4}\,h^3\,\mathrm{Mpc}^{-3}]$	$b(z)$		z	$n(z)$ $[10^{-4}\,h^3\,\mathrm{Mpc}^{-3}]$	$b(z)$
2.0	25	2.5		4.0	1.5	5.8
2.5	12	3.3		4.5	0.8	6.6
3.0	6.0	4.1		5.0	0.4	7.4
3.5	3.0	4.9				

Table 1: Our 'idealised' sample: a $m_{UV} = 24.5$ magnitude-limited dropout sample as defined by Ref. [10]. Here $n(z)$ and $b(z)$ correspond to the expected number density and linear galaxy bias with redshift.

z	$n(z)$ $[10^{-4}\,h^3\,\mathrm{Mpc}^{-3}]$	$b(z)$		z	$n(z)$ $[10^{-4}\,h^3\,\mathrm{Mpc}^{-3}]$	$b(z)$
2.0	9.8	2.5		4.0	1.0	3.5
3.0	1.2	4.0		5.0	0.4	5.5

Table 2: Our 'fiducial' sample achievable with next generation facilities. The number density and galaxy bias estimates derive from Refs. [10, 30, 31, 32, 33] and [34]. Note this is significantly less dense than that in Table 1 at lower redshift; We find the limiting factors are efficient pre-selection of LAEs based on broad-band imaging, LSST u-band depth and our posited survey speed for $z = 2, 3$ and 4 respectively.

2 Forecasts

2.1 Primordial non-Gaussianity

We follow Ref. [13] in order to forecast the constraints on primordial non-Gaussianity achievable with these samples. The results are shown in Table 3 when including both the power spectrum and bispectrum. We find that local f_{NL} sees the largest improvement, achieving $\sigma(f_{NL}^{\mathrm{local}}) \approx 0.1$ for the fiducial sample. This represents a factor of $\simeq 50$ improvement over current surveys and achieves the precision necessary for a paradigm shift in our understanding of the early Universe. No planned survey can deliver this at such a redshift, which would be entirely complementary to lower z studies [35]. When including the external CMB and LSS data expected to be available by first light, the constraints on equilateral and orthogonal f_{NL}^{local} see additional improvements of ~ 2 and 3 over current estimates. Given this achievable precision, the measurement will likely be systematic dominated and the survey should be designed accordingly.

The competitiveness of spectroscopy is clear from the sharp degradation in constraints – a factor of 3 for both local and orthogonal, and a factor of 4 for equilateral – if only photometric redshifts are available.

2.2 Dark Energy

The galaxy power spectrum yields measures of the expansion and growth rates. In turn, these can be used to infer the energy content at a particular redshift. In Figure 1, we show

$\sigma(f_{NL})$ Fiducial / Idealised	P	$+B$	+ External	Current (Planck)	Photo-z degradation
Local	0.75 / 0.63	0.11 / 0.073	0.11 / 0.073	5	×3
Equilateral	–	43 / 23	23 / 18	43	×4
Orthogonal	50 / 33	8.8 / 5.0	7.5 / 4.7	21	×3

Table 3: Constraints on f_{NL} for the two samples considered. P denotes those derived from the power spectrum, while $+B$ includes additional constraints from the bispectrum. External datasets include constraints on f_{NL} coming from Planck [36], DESI [37] and Simons Observatory [4], which are expected to complete by our first light. In the last column, we illustrate a photo-z degradation corresponding to $\sigma(z)/(1+z) = 2 \times 10^{-2}$.

that both potential surveys constrain the fraction of Dark Energy to one percent, or even sub-percent, precision to $z \sim 5$. This would represent a tremendous increase in precision over DESI, especially for $z > 3$. In the standard parametrization, these correspond to a Dark Energy Figure of Merit (FoM) of 398 and 441 for the fiducial and idealised samples respectively. This is an improvement of a factor of 2.7 over DESI [37] when combined with the current Planck constraints. Spectroscopy is essential in this respect, with a degradation of over $\sim 60\%$ for photometric redshifts ($\sigma(z)/(1+z) = 0.01$).

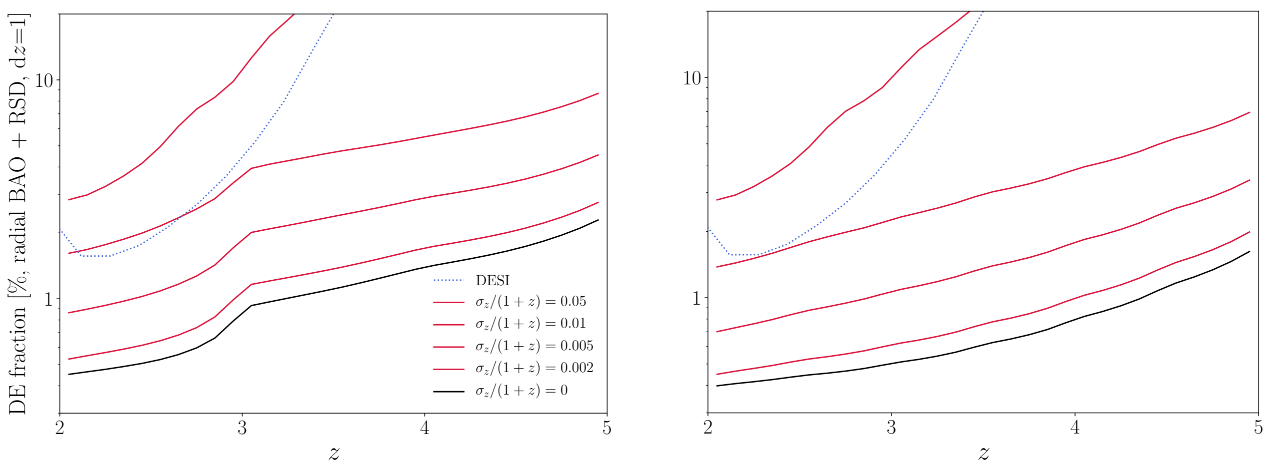

Figure 1: The absolute error on the fraction of Dark Energy, Ω_{DE}, at a given redshift for the fiducial (left) and idealised (right) samples. This is obtained from a combination of radial Baryon Acoustic Oscillation (BAO) and Redshift-Space Distortions (RSD). If Dark Energy is a cosmological constant, its fraction is forecasted to be 7%, 3%, 2% and 1% at $z = 2, 3, 4, 5$ to a very high degree of accuracy, which motivates facilities capable of challenging this prediction.

Table 4 shows forecasts for the (beyond) Standard Model parameters. In addition to the Dark Energy FoM, large improvements are found for the curvature, Ω_K, (with errors decreasing by over a factor of 2), together with the sum of neutrino masses.

While not explored in great detail here, it has been shown that cross-correlation with the CMB and Intensity mapping experiments can greatly reduce systematics and break several astrophysical and cosmological degeneracies. As an example, Figure 2 shows constraints on the amplitude of fluctuations $\sigma_8(z)$ as a function of redshift by cross-correlating CMB lensing with galaxy surveys. With this potential for synergy with future CMB surveys, we

Parameter	σ(parameter) Fid./Ideal.	DESI
Curvature $\Omega_K/10^{-4}$	6.6 / 5.2	12.0
Neutrinos $\sum m_\nu$	0.028 / 0.026	0.032
Spectral index n_s	0.0026 / 0.0026	0.0029
Running α_s	0.003 / 0.003	0.004
Rel. species N_{eff}	0.069 / 0.069	0.078
Gravitational slip	0.008 / 0.008	0.01
D.E. FoM	398 / 441	162

Table 4: Forecasts on cosmological parameters from our samples, combined with Planck priors. Gravitational slip is defined as the ratio between the two potentials describing the metric, in combination with a CMB experiment with map noise of 1 μK-arcmin.

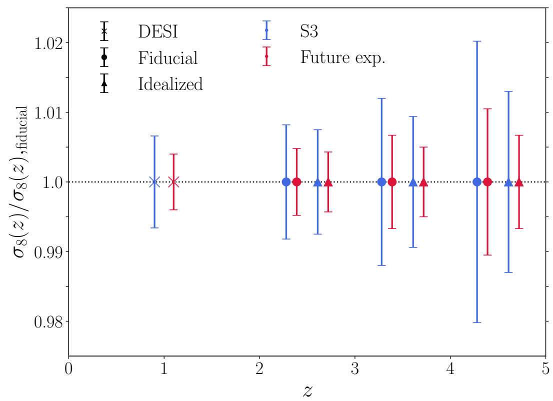

Figure 2: Constraints on $\sigma_8(z)$ from cross-correlation with CMB lensing. 'S3' and 'Future exp.' refer to CMB experiments with map noise of 7 and 1 μK-arcmin respectively.

can extract sub-percent constraints on the growth that are relatively insensitive to the $z < 2$ universe and hence a powerful probe of non-standard Physics.

3 Challenges

Further development of efficient pre-selection of LAEs from broad-band photometry is a requirement for this case as presented. The success of this pre-selection will largely determine the necessary facilities and achievable samples. Some of the measurements outlined above – especially local f_{NL} – also require complete understanding of e.g. the parent photometry and the galaxy selection function generally [2, 38, 39]. Percent-level sky subtraction with fibers and exposures approaching an hour, together with mitigation of line confusion, are also technical tasks to be overcome. Potential strategies have already been proposed and are under active study, but future surveys will require careful consideration of these points during any design phase.

4 Conclusions

The colossal, relatively uncharted, volume at $z > 2$ and known means of efficiently selecting high-z galaxies grants a tremendous opportunity to study the beginning and fate of our Universe, namely Inflation and Dark Energy. We have shown potential surveys can test the early Universe (Gaussianity) up to a factor of ~ 50 better than our current bounds and cross the highly significant threshold of $f_{NL} \simeq 1$ that would separate single-field from multi-field models of Inflation; Such measurements would be entirely complementary to low-z studies. This is enabled by spectroscopic redshift precision, with photometric redshift precision degrading these constraints by a factor of three or greater.

Such a dataset would leave an important legacy for the science cases we have presented, together with a wealth of opportunities for the field of galaxy formation and many others.

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